

A SEARCH FOR ULTRAVIOLET-EXCESS OBJECTS

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Abstract

A search for UV-excess objects has been undertaken with the use of the 105-cm Schmidt telescope at the Kiso Observatory by means of UGR three-image method. A general description of the method is given, and positions, magnitudes, and color indices of detected 588 objects covering 22 fields (600 square degrees) are listed in this paper. The brightness of the objects ranges from 12.5 to 18.5 magnitudes. About 13 percent of all are identified with white dwarfs or quasars appeared in the previous catalogues, and the remainder is left unidentified.

Key words: Color survey; UV-excess objects; Schmidt observation.

1. Introduction

A number of color surveys followed the pioneering work by Humason and Zwicky (1947) for the detection of blue stars including white dwarfs, subdwarfs, novae, quasars, compact galaxies, and some other peculiar objects. Proper motion surveys by Luyten (1951) and Giclas (1960) were also efficient for the detection of white dwarfs.

Recently a color survey was done by Green (1976) covering about 10,000 square degrees of $|b| \geq 30^\circ$. He discovered four quasars, eighty-nine white dwarfs, and many other blue objects brighter than 16.5 mag (Green and Morrill 1978, Green 1980*). Berger and Fringant (1977*, 1980*) obtained three-image plates with the Palomar 48-inch Schmidt telescope. They picked out over 7,000 blue objects down to 18.5 or 19.0 mag in 23 fields near the north and the south galactic poles. Steppe (1978*) made a systematic search for blue objects in high galactic latitudes. He used several plates per color band taken with the Palomar 48-inch Schmidt telescope. A total of 1,906 objects are listed down to 20 mag in five high latitude fields.

We have undertaken a search for UV-excess objects with the use of the 105-cm Schmidt telescope at Kiso Observatory. The observation technique is a kind of "three-image method" (e.g., Haro and Luyten 1962) by which we obtain three images for individual objects on a plate through *U* (ultraviolet), *G* (green), and *R* (red) filters. This method is very profitable to detect UV-excess objects quickly by visual inspection.

2. Observation

The 105-cm Schmidt telescope has a UBK7 corrector plate which is transparent in the ultraviolet band. This telescope also has a mechanism of the automatic filter loading in its tube. This is very efficient for the three-image method because three exposures can be done successively without changing the attitude of the telescope (Takase et al. 1977). The plate scale is 62.6 mm^{-1} , and a 14 inch square plate covers about $6^\circ \times 6^\circ$ celestial area. Kodak 103aE emulsion (baked at 40°C during

The asterisk * is attached to the literature cited in table 2.

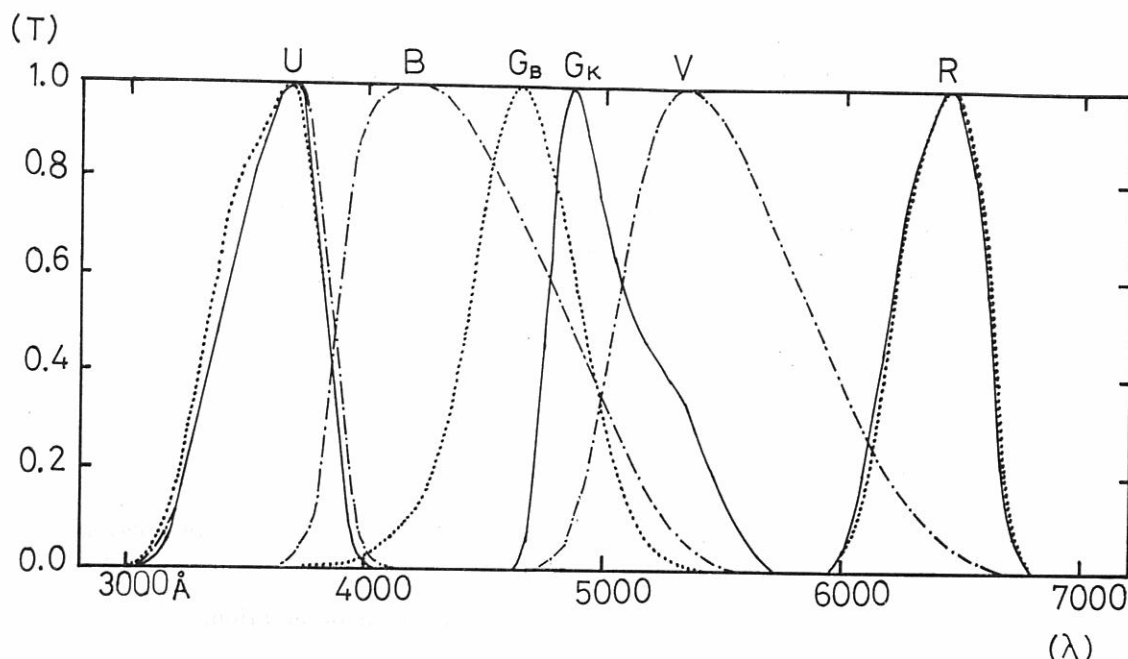


Fig. 1. Relative response curves of Kiso UGR system (full lines), Becker's RGU system (dotted lines), and Johnson's UB system (dot-dash lines). Curves for two latter systems are referred to Buser (1979).

17 hours in forming gas) are used behind filters UG-1, BPB-50, and RG-610; the second one is supplied by Fuji Film Corporation, and the other two are Schott color filters.

Relative response curves of this color system are shown in figure 1, together with Johnson's UB and Becker's RGU systems (Buser 1978). Our *U* band closely resembles those of Johnson's and Becker's. Our *R* band is the same combination of emulsion and filter as Becker's *R*. Our *G* band has a unique sensitivity; it is located between *B* and *V*, and shifted to the longer wavelength by 220 Å than Becker's *G*.

U, *G*, and *R* images are taken in alignment with the separation of 20 arcseconds each other on the plate (cf., figure 2). The standard exposure times are 40, 100, and 20 minutes for *U*, *G*, and *R* bands, which give almost the equal size or density of the three images for unreddened A0 stars. In the early period of the survey, different combinations of emulsion like 103aF or 103aE without baking and exposure times were applied to some observations.

The main object area is the region from the north to the south galactic pole with the galactic longitude $l^{II}=180^\circ$. A total of about 100 Kiso fields are included in the survey. Table 1 lists the fields and plates searched in this paper. The first column designates the Kiso field number (cf., *Kiso Information Bulletin*, vol. 1, 1979), and the following four columns give the coordinates of the plate center. The last five columns include plate data.

3. Detection and Measurement

The detection has been made by visual inspection with the use of the projector whose field is a diameter of 16 mm on the photographic plate with five times magnification. The searching time over a plate is two to three hours in the high galactic latitudes, and more in the field near the galactic plane. A plate is searched repeatedly by each of us in order to secure as complete detection as possible. Initially we intended to detect bright objects in $(U+R-2G)$ magnitude. But now, Kiso

Table 1

Observed Areas and Plate Data

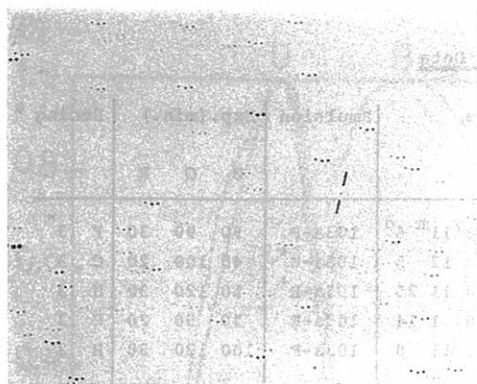
Area	Plate Center Coordinates (1950)				Plate No.	Date	Emulsion	Exp. (min.)			Seeing *
	α	δ	z^{II}	b^{II}				U	G	R	
1037	2 ^h 00 ^m	-10° 00'	170.26	-65.87	KL1816	1978 Y11 ^m 6 ^d	103a-F	90	90	30	F 3"
896	3 00	0 00	177.42	-48.30	1895	1978 12 5	103a-E ⁺	40	100	20	C 3
897	3 20	0 00	182.27	-44.59	1865	1978 11 25	103a-E ⁺	50	120	30	H 3
684	4 20	15 00	180.02	-23.58	1330	1978 1 14	103a-E	50	50	20	C 3
685	4 40	15 00	183.25	-19.79	1825	1978 11 8	103a-F	100	120	30	H 3
757	4 40	10 00	187.57	-22.75	1862	1978 11 24	103a-E ⁺	40	100	25	C 3
686	5 00	15 00	186.22	-15.88	1817	1978 11 6	103a-F	90	-	30	C 3
543	5 20	25 00	180.56	- 6.33	1818	1978 11 7	103a-F	90	92	30	C 3
476	7 00	30 00	186.85	15.49	1935	1978 12 25	103a-E ⁺	40	100	20	C 3
343	8 00	40 00	180.59	30.21	1557	1978 4 4	103a-E	60	120	25	C 5
344	8 24	40 00	181.56	34.73	1554	1978 3 31	103a-E	60	120	25	C 4
345	8 48	40 00	182.25	39.29	1445	1978 2 26	103a-E	50	90	20	C 3
410	9 00	35 00	189.10	41.10	1446	1978 2 26	103a-E	25	50	10	C 5
347	9 36	40 00	182.55	48.48	1947	1978 12 26	103a-E ⁺	40	100	20	C 3
348	10 00	40 00	181.98	53.06	1897	1978 12 5	103a-E ⁺	40	90	20	C 3
352	11 36	40 00	168.48	70.34	1519	1978 3 17	103a-E	45	80	15	C 1
494	13 00	30 00	80.80	86.45	1555	1978 3 31	103a-E	60	120	25	F 3
433	16 40	35 00	56.97	40.70	2139	1979 3 28	103a-E ⁺	40	92	20	H 2
4109	18 17	66 40	96.59	28.13	2254	1979 6 24	103a-E ⁺	36	90	18	F 2
49	21 00	75 00	109.72	18.68	1716	1978 10 6	103a-F	90	60	30	F 3
8741	23 13	13 20	90.18	-43.08	1717	1978 10 6	103a-F	90	60	30	C 2
813	23 20	10 00	89.97	-46.83	1860	1978 11 22	103a-E	60	150	30	F 3

* Characters represent the condition of the sky: C is clear, F is fair, and H is hazy. Numerals express the size of seeing image roughly.

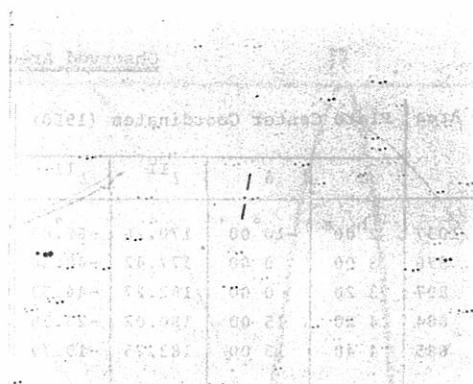
UV-excess objects (written as KUV objects abbreviately in the following) are detected with a criterion of enhanced U image relative to G and R images, because we have experienced that stressing on both U and R requires much time of search.

Candidates are carefully inspected with a magnifier, and the eye-estimated G magnitude and the color index are assigned to them. To accomplish this process, a sequence of the standard object of magnitude must be set on individual plates. There exist few fields in which objects are photometrically observed as faint as 18 or 19 magnitudes. In this circumstance, V magnitudes of bright standards are referred to Purgathofer (1969), Mermilliod (1976), and Mermilliod and Nicolet (1977). In the portion of fainter magnitudes, objects are picked out of the lists of Eggen and Greenstein (1965*, 1967*), and Burbidge, Crowne, and Smith (1977*). Moreover, the *Palomar Sky Survey Print* is consulted in order to check the identification of objects and magnitudes. The magnitude thus obtained is the brightness of G band calibrated with V magnitude.

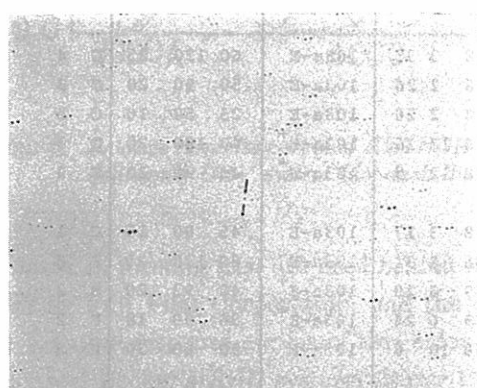
The color index $C.I.$ is estimated as the enhancement in U image relative to G image, and calibrated with the known values of color $U-B$, which are again referred to the papers cited in the last paragraph. It is determined with a step scale of every 0.5. The value $C.I. = -3$ corresponds to



(a) 16484+3706



(c) 16454+3234



(b) 16501+3404



(d) 16376+3331

Fig. 2. Portions of a searched plate KL2139 of the field A0433. Every star has U , G , and R images (from left to right) separated by about 20 arcseconds. (a) $15.^m8$, $C.I. = -0.5$. (b) $14.^m5$, $C.I. = -0.5$. (c) EG239 (DB); $V = 13.65$, $B - V = -0.11$, $U - B = -1.04$, $C.I. = -2.0$. (d) EG120 (DAs); $V = 14.64$, $B - V = +0.22$, $U - B = -0.58$, $C.I. = -1.0$.

$U - B = -1.5$, and $C.I. = 0.0$ to $U - B = 0.0$. The accuracy of magnitude is about ± 0.5 mag in the well-calibrated field, though it is worse in the poorly-calibrated field. This inaccuracy partly comes from the error in eye-estimation of the magnitude, and partly comes from the deficiency of photometric data.

Figure 2 shows four portions of the plate KL2139 of the field A0433 ($\alpha_{1950} = 16^h 40^m$, $\delta_{1950} = +35^\circ$). Every object has U , G , and R images from left to right. Four examples show varieties of U enhancement. Among them, the objects (c) and (d) belong to Eggen and Greenstein's.

In order to examine our color index, the iris measurement was made on the plate KL1445 of the field A0345 ($\alpha_{1950} = 8^h 48^m$, $\delta_{1950} = +40^\circ$). The result is illustrated in figure 3, where the coordinates are the differences between two iris values, $I_U - I_G$, and $I_G - I_R$. Since the iris value adopted here holds a good linearity to the magnitude over a wide range of it (Ishida, Maehara, and Ohashi 1978), it is noticed that the $C.I.$ is well correlated with the color $U - G$ for UV-excess objects. Two labelled objects GD94 and GD98 are a suspected white dwarf and a white dwarf, respectively.

The position measurement of the objects is performed with the use of a comparator at Kiso. Standard stars are selected at density of a star per square degree from AGK3 catalogue for the field

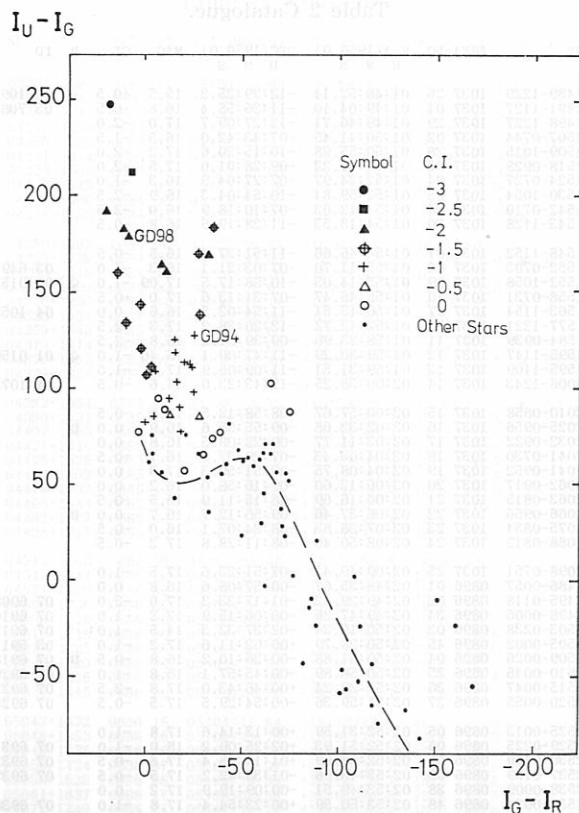


Fig. 3. Iris measurements on plate KL1445 of the field A0345. Both coordinates are the differences in iris values of respective bands. KUV objects are plotted as the symbols corresponding to their values of C.I. Two labelled objects GD94 and GD98 are white dwarfs whose $U-B$'s are -0.60 and -0.94 , respectively.

of $\delta \geq -2^\circ$, and from SAO catalogue for the southern field. The position of each standard star is measured three times, and those of KUV objects are measured successively on the plate. The reduction of the position is carried out by the standard coordinate method (Mikami 1980). The positional accuracy is within about $\pm 0.''5$ (s.d.) for most stars, though it is worse for stars located near the edge of the plate.

4. Catalogue of the KUV objects

Data of individual KUV objects are listed in table 2. The first column shows the KUV designation, which is composed of the right ascension and declination of the objects. The second column gives the Kiso field number and the serial number in the field. The following two columns give the right ascension and declination of the object at the equinox of 1950.0. The fifth and the sixth columns present the visual G magnitude and the $C.I.$ In the seventh column the remark is given as follows:

- W: white dwarfs,
- Q: quasars,
- D: diffuse objects,
- V: variable objects.

The last column gives the identification of stars previously appeared in the literature; the number

Appendix of Table 2

- Ref. 01; Quasars
Burbidge, G.R., Crowne, A.H., and Smith, H.E. 1977, *Astrophys. J. Suppl.*, **33**, 113.
- Ref. 02; EG stars
Eggen, O.J., and Greenstein, J.L. 1965, *Astrophys. J.*, **141**, 83.
Eggen, O.J., and Greenstein, J.L. 1965, *Astrophys. J.*, **142**, 925.
Eggen, O.J., and Greenstein, J.L. 1967, *Astrophys. J.*, **150**, 927.
Greenstein, J.L. 1969, *Astrophys. J.*, **158**, 281.
Greenstein, J.L. 1970, *Astrophys. J. (Letters)*, **162**, L55.
Greenstein, J.L. 1974, *Astrophys. J. (Letters)*, **189**, L131.
Greenstein, J.L. 1975, *Astrophys. J. (Letters)*, **196**, L117.
Greenstein, J.L. 1976, *Astrophys. J. (Letters)*, **207**, L118.
- Ref. 03; LP stars
Luyten, W.J. 1979, "White Dwarfs II", Univ. Minnesota, Minneapolis.
- Ref. 04; GD stars
Giclas, H.L., Burnham, R. Jr., and Thomas, N.G. 1965, *Lowell Obs. Bull.*, **VI**, 155.
Giclas, H.L., Burnham, R. Jr., and Thomas, N.G. 1967, *ibid*, **VII**, 49.
Giclas, H.L., Burnham, R. Jr., and Thomas, N.G. 1970, *ibid*, **VII**, 183.
Giclas, H.L., Burnham, R. Jr., and Thomas, N.G. 1972, *ibid*, **VII**, 217.
Giclas, H.L., Burnham, R. Jr., and Thomas, N.G. 1973, *ibid*, **VII**, 273.
Giclas, H.L., Burnham, R. Jr., and Thomas, N.G. 1975, *ibid*, **VIII**, 9.
- Ref. 05
Green, R.F. 1980, *Astrophys. J.*, **233**, 685.
- Ref. 06
Steppe, H. 1978, *Astron. Astrophys. Suppl.*, **31**, 209.
- Ref. 07
Berger, J., and Fringant, A.M. 1977, *Astron. Astrophys. Suppl.*, **28**, 123.
Berger, J., and Fringant, A.M. 1980, *Astron. Astrophys. Suppl.*, **39**, 39.

of first two figures indicates the reference number appended at the end of the table, and the number of following figures indicates the object number in the literature.

5. Discussion

The three-image method was initially adopted by Haro and Herbig (1955). It is very efficient to detect blue objects. Their work was followed by Iriarte and Chavira (1957), Haro and Luyten (1962), Rubin, Moore, and Bertiau (1967), and Berger and Fringant (1977*, 1980*). All of these observations employed the *UBV* system. The *UGR* system adopted here is more sensitive to the gradient of continuous spectra of the objects. This fact is certified by the comparison of the *UBV* three-image plate with the *UGR* plate. Takase (1980) has detected some 1,100 galaxies with bright ultraviolet radiation by means of the similar method as ours. He picked out diffuse objects on the *UGR* three-image plates, while our search is mainly confined to stellar images with UV-excess. Six out of twenty plates are common to our search.

This work is compared with the following surveys. The comparison with Green's (1976) search cannot be done, since lists of individual detected objects are not available for us. On the other hand, both Steppe (1978*), and Berger and Fringant (1977*) listed the objects around the north galactic pole. Photometric values by Steppe (1978*) are derived from several plates in each band, and they are considered to be the most reliable ones. In figure 4, our color indices are plotted against Steppe's $U-B$ for common objects. Though the scatter is fairly large especially in the portion of small UV-excess, the $C.I.$ is as a whole proportional to the color $U-B$ with the inclination of about 2.5. This relation holds over a range of four magnitudes. Therefore, the eye-estimated color index is proved to be a good indicator of the degree of the UV-excess. On the other hand, our color

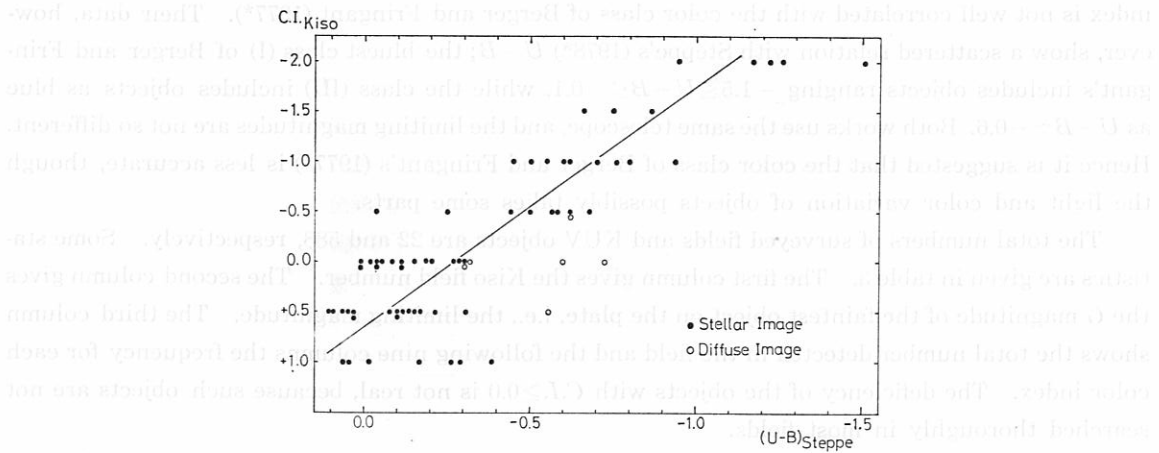


Fig. 4. Comparison of our color index $C.I.$ with $U-B$ by *Steppe* (1978*) in the region including SA57 (plate KL1555).

Table 3
Limiting Magnitudes and Number of KUV-Objects 22 Fields

Area	Limiting Magnitude	Total Number	Color Indices								
			-3.0	-2.5	-2.0	-1.5	-1.0	-0.5	0	0.5	1.0
1037	17.5	31	0	4	3	1	6	7	6	4	0
896	18.0	52	0	5	2	3	10	14	13	5	0
897	18.2	20	0	1	3	3	2	8	1	2	0
684	17.2	10	0	1	1	2	3	2	1	0	0
685	17.0	14	0	4	2	2	2	2	2	0	0
757	17.2	3	0	0	1	1	1	0	0	0	0
686	17.5	24	0	2	4	4	4	6	4	0	0
543	17.2	22	0	1	0	1	3	6	10	1	0
476	18.0	18	0	1	4	4	3	4	2	0	0
343	17.8	34	0	1	4	6	6	8	7	2	0
344	17.8	24	0	0	3	5	8	4	3	0	1
345	18.0	67	1	1	5	7	14	10	23	4	2
410	16.8	9	0	0	0	2	1	6	0	0	0
347	18.5	48	0	0	5	12	8	8	10	5	0
348	18.2	22	0	1	1	6	6	3	4	1	0
352	17.5	26	0	1	4	3	6	4	7	1	0
494	18.0	72	0	0	4	11	13	15	16	11	2
433	17.2	17	0	2	7	0	4	3	0	1	0
4109	17.2	17	0	1	4	5	4	2	1	0	0
49	17.0	8	0	1	0	2	2	2	1	0	0
8741	17.2	16	0	0	1	3	3	5	4	0	0
813	17.6	34	0	0	3	7	9	7	7	1	0
		588	1	27	61	90	118	126	122	38	5

index is not well correlated with the color class of Berger and Fringant (1977*). Their data, however, show a scattered relation with Steppe's (1978*) $U-B$; the bluest class (I) of Berger and Fringant's includes objects ranging $-1.5 \leq U-B \leq -0.1$, while the class (III) includes objects as blue as $U-B \approx -0.6$. Both works use the same telescope, and the limiting magnitudes are not so different. Hence it is suggested that the color class of Berger and Fringant's (1977*) is less accurate, though the light and color variation of objects possibly takes some parts.

The total numbers of surveyed fields and KUV objects are 22 and 588, respectively. Some statistics are given in table 3. The first column gives the Kiso field number. The second column gives the G magnitude of the faintest object on the plate, i.e., the limiting magnitude. The third column shows the total number detected in the field and the following nine columns the frequency for each color index. The deficiency of the objects with $C.I. \geq 0.0$ is not real, because such objects are not searched thoroughly in most fields.

The integrated luminosity function of the KUV objects are shown in figure 5, where they are divided into three groups according to the $C.I.$ The completeness in detection is not secured for the objects fainter than 17 mag, because the limiting magnitude of underexposed plates are located around it. As is shown in this figure, objects with the large UV-excess are more frequently detected around 15 mag than expected from the uniform distribution. On the other hand, objects with the small UV-excess do not exhibit such a hump. This tendency is clearly seen in the fields A0684, A0685 (Hyades), and A0433 (M13 region). Some of them are identified as white dwarfs or blue horizontal-branch stars in previous catalogues.

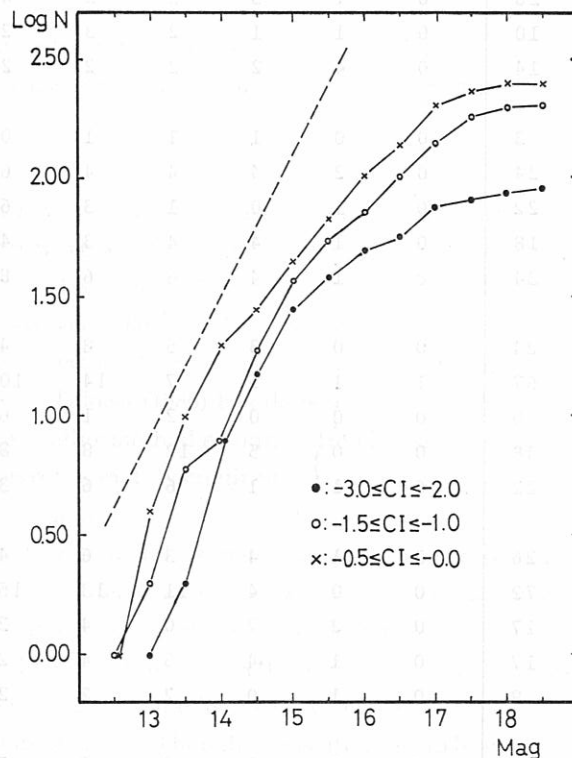


Fig. 5. Logarithmic integrated luminosity function of the detected KUV objects. They are divided into three groups according to $C.I.$ The completeness is not secured at the fainter portion than 17 mag. The dashed line indicates uniform distribution of objects.

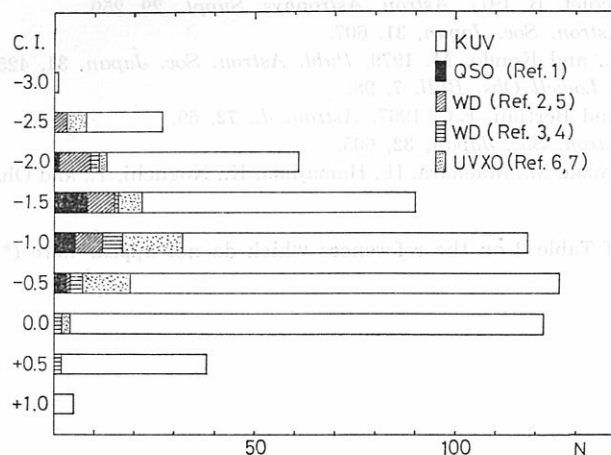


Fig. 6. Frequency distribution against *C.I.* for the detected KUV objects. Common objects with other surveys are designated with hatches.

The frequency distribution against the *C.I.* for all the objects of this survey is shown in figure 6. 17 Objects are identified as quasars, 26 as spectroscopic white dwarfs, and 15 as suspected white dwarfs. Among the remaining objects, 42 are identified as blue objects of Berger and Fringant (1977*, 1980*) and Steppe (1978*). The remainder (83 percent) is left unidentified. A suspected dwarf nova has been discovered in the course of this search (Noguchi et al. 1979). The *C.I.*'s of quasar range from -0.5 to -2.0 , while those of white dwarfs are distributed down to -2.5 .

Judging from the result, major constituents of our objects are possibly white dwarfs. Quasars and blue extragalactic objects seem to be minor groups in our accessible magnitudes. Moreover, galactic objects such as hot subdwarfs, novae, dwarf novae, central stars of planetary nebulae, etc. are also expected as candidates (Green 1980*). Spectroscopic observations are undertaken with the use of the 188-cm reflector at Okayama Astrophysical Observatory.

We are grateful to Professors B. Takase, Y. Yamashita, and K. Ishida for reading the manuscript, and for giving valuable advice. We also thank the staff of the Kiso Observatory for their assistance in observations. Prof. Takase kindly permitted us to add his plates in this search. The computation was performed on the OKITAC-50 at Kiso and the FACOM 230 at Mitaka. Drs. K. Hurukawa, S. Okamura, T. Mikami, and K. Hamajima kindly permitted us to use their computation programs.

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Mikami, T. 1980, *Publ. Astron. Soc. Japan*, **31**, 607.
Noguchi, T., Maehara, H., and Kondo, M. 1979, *Publ. Astron. Soc. Japan*, **31**, 425.
Purgathofer, A. Th. 1969, *Lowell Obs. Bull.* **7**, 98.
Rubin, V.C., Moore, S., and Bertiau, E.C. 1967, *Astron. J.*, **72**, 59.
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Takase, B., Ishida, K., Shimizu, M., Maehara, H., Hamajima, K., Noguchi, T., and Ohashi, M. 1977, *Ann. Tokyo Astron. Obs.*, **16**, 74.

N.B., See the appendix of Table 2 on the references which do not appear here (* in the text).